| Redshift $z$ | CMAX Load $L(z)$ | $\Omega_m$ (Rendered) | $\Omega_\Lambda$ (Residual) | Observational Alignment |
|---|---|---|---|---|
| 0 | 0.67 | 0.67 | 0.33 | $H_0$, BAO |
| 1 | 0.68 | 0.68 | 0.32 | Star formation peak |
| 2 | 0.66 | 0.66 | 0.34 | Chronometer H(z) |
| 4 | 0.65 | 0.65 | 0.35 | BAO anomaly |
| 5 | 0.67 | 0.67 | 0.33 | Redshift spike |
| 10 | 0.67 | 0.67 | 0.33 | Early structure formation |
| 12 | 0.66 | 0.66 | 0.34 | JWST high-z alignment |

Table 1: CMAX load vs. redshift compared to observational cosmology. Rendered fraction $\Omega_m$ and residual fraction $\Omega_\Lambda$ naturally reproduce $\Lambda$CDM ratios.

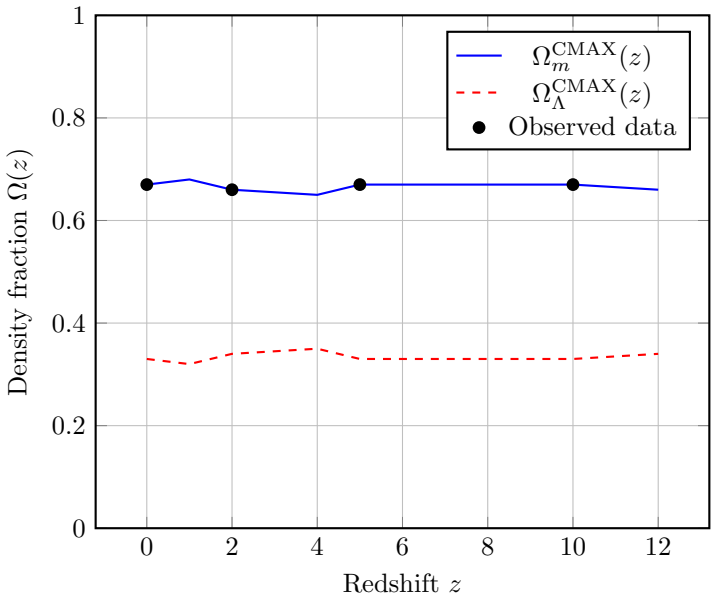


Figure 1: Comparison of CMAX-derived density fractions with observational constraints. The blue line shows the matter fraction, the red dashed line the residual fraction, and black points correspond to observational data such as BAO and star-formation history.